ver, in this paper we tried to describe the works on detailization of the evolutionary directions of separate types of the eclipsing binaries and on interrelations between these types. It seems to us, that the shown results confirm:

1) the idea of the transitions from one type of the eclipsing binary stars into another;

2) the binarity of the "contact" evolutionary stage of the systems, at which a part of the stars finishes its way as the binary stars (low–mass, short–period $MS-$, $SimCW-$, $CW-$, $CE-$ systems of stars), merging into one object; the rest evolving to other types of the eclipsing binary stars with the formation of subgiant stars.

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ON ESTIMATES OF LOWER AND UPPER LIMITS FOR THE MASSES OF COMPACT COMPONENTS IN CLOSE BINARIES

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ABSTRACT. A method of the determination of limits for a compact component mass on the base of disk emission lines parameters is described. Lower limit of mass depends upon the distance between maxima in double peaked lines, upper upon full width of line. The method is tested for some cataclysmic variables with well–known masses of compact components. It is obtained a lower limit for the mass of the compact object in the close binary SS 433 is $4.9 M_\odot$. This component apparently is a black hole.

Key words: Stars: Binaries, Accretion Disks